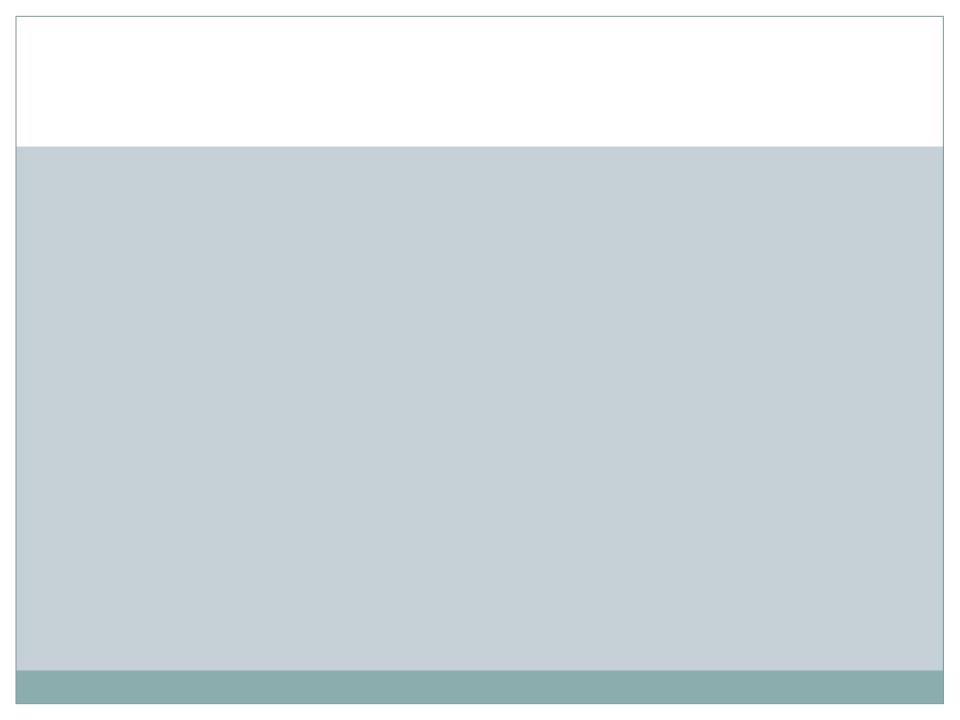
Computational Modeling Decay Spectroscopy

40 , 00

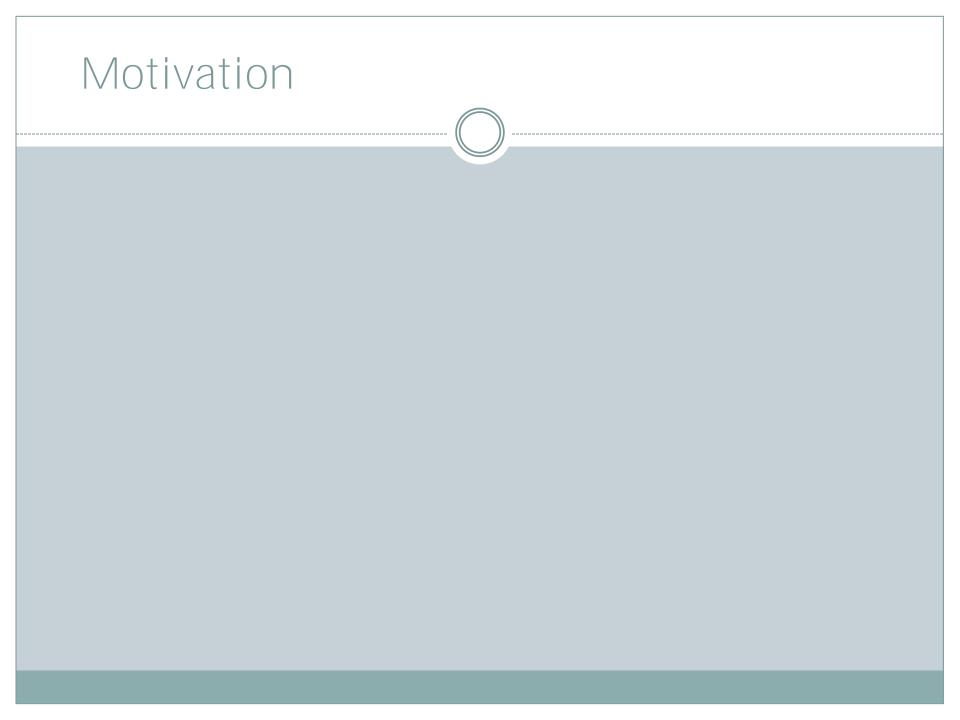






Motivation Solar Reactions $p(\eta p)$ HTV_e 10⁻⁵ % \rightarrow ³He+n⁺ \rightarrow ⁴He $^{2}\text{H}+\text{n}^{+}\rightarrow ^{3}\text{He}+\gamma$ $e + e^+ + v$ 7月暮日。 VIBUSEN 3110 /110 >730 y $^7\mathrm{Be+p^+} \rightarrow ^8\mathrm{B} + \gamma$ $pp\Pi$ ppl 8Da* \4Ua\4Ua ррШ









Physics Package Used

Electron Gamma Shower national research council (Canada) (EGSnrc)



Monte Carlo Techniques (2)

Each particle has a chance of undergoing various interactions

Pair production, Compton scattering, etc.

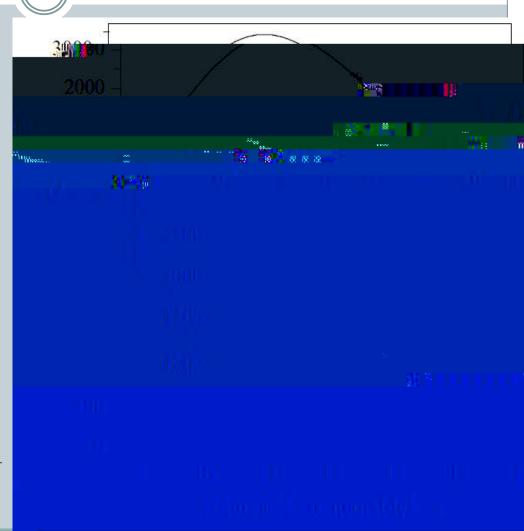
A random number is generated and corresponds to an outcome

Any extra particles created are also followed to the cut off energy To get high accuracy we run lots of particles

Our simulations require 10⁷ to 10⁸ source particles

Past Work

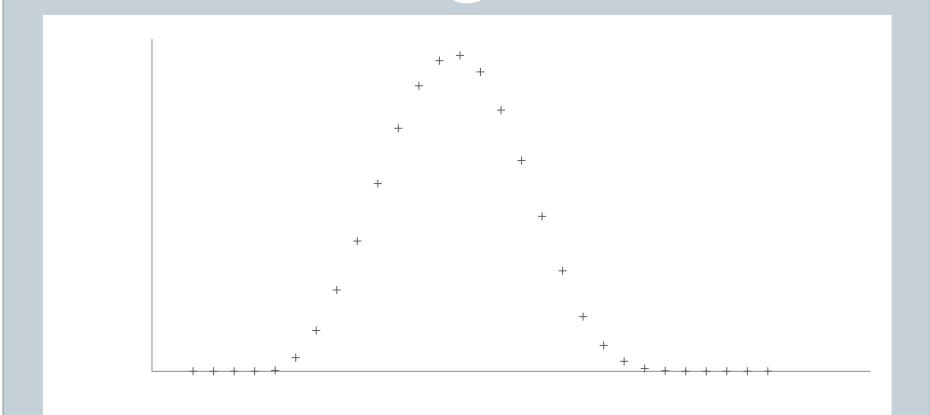
Past Measurement



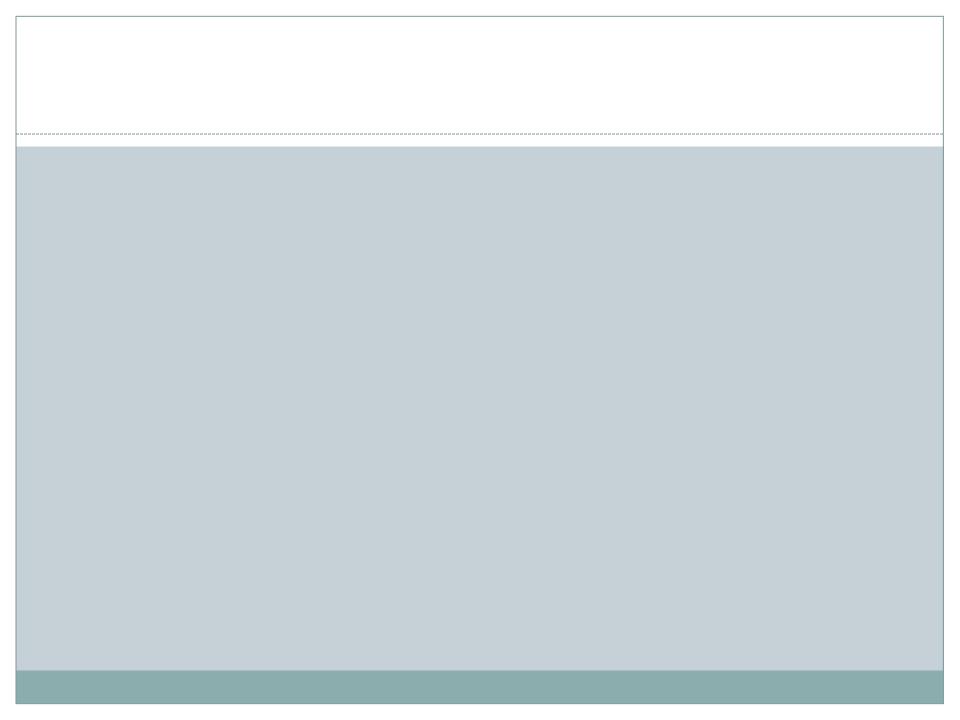
W.T. Winter



EGS Modeling

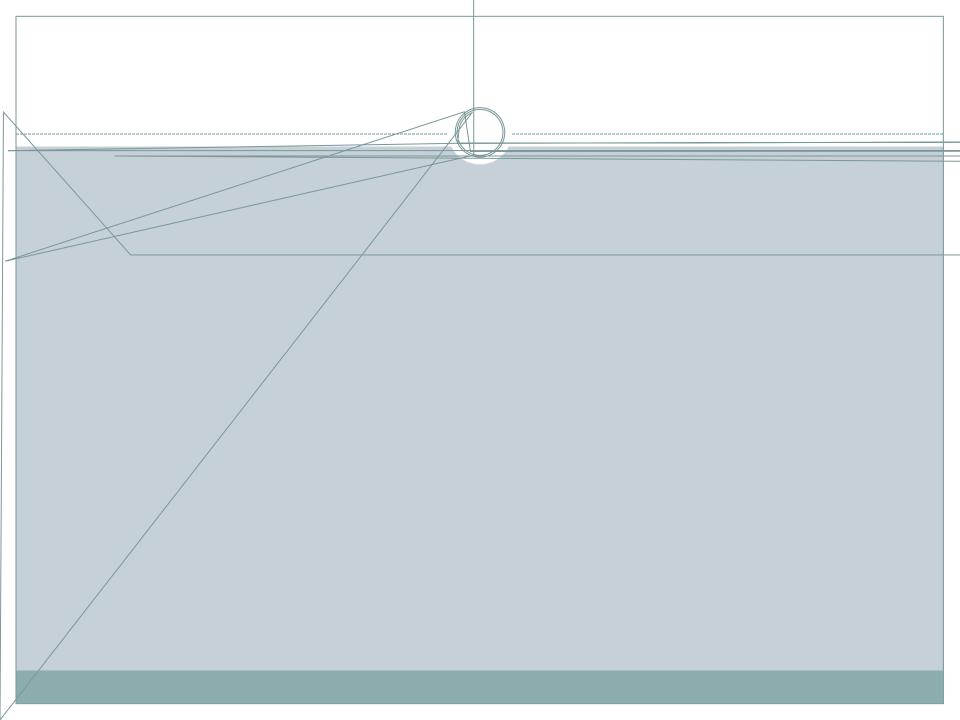


Response Curve at 12 MeV/c









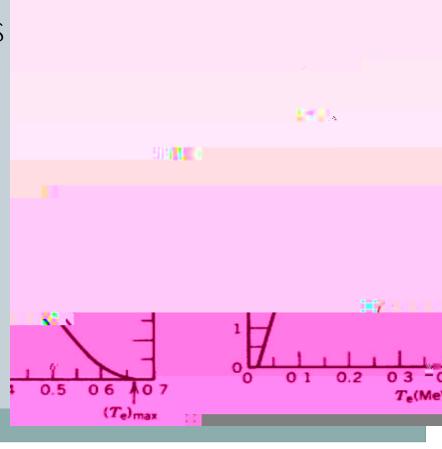
The Fermi Function

Based on Coulomb interaction between emitted particle and positive nucleus

$$F(Z, T_e) = \int_{A} A + \frac{B}{T_e - 1}$$

$$A = mZ + K$$

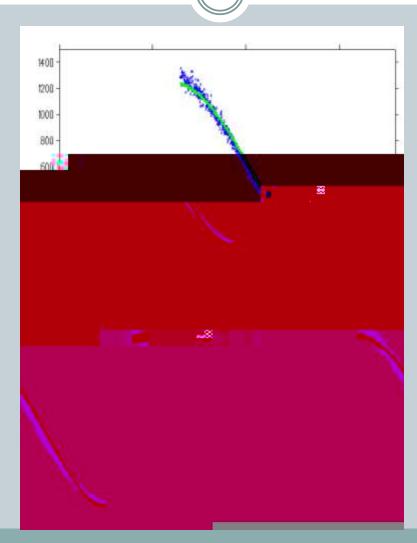
$$B = aZ e^{bZ}$$



Comparison of Fits

With Fermi:

Allowed:



Y-Axis: Counts

X-Axis: MeV

